

Diving below the spin-down limit: Constraints on gravitational waves from the energetic young pulsar PSR J0537–6910

- R. ABBOTT,¹ T. D. ABBOTT,² S. ABRAHAM,³ F. ACERNESE,^{4,5} K. ACKLEY,⁶ A. ADAMS,⁷ C. ADAMS,⁸ R. X. ADHIKARI,¹ V. B. ADYA,⁹ C. AFFELDT,^{10,11} D. AGARWAL,³ M. AGATHOS,^{12,13} K. AGATSUMA,¹⁴ N. AGGARWAL,¹⁵ O. D. AGUIAR,¹⁶ L. AIELLO,^{17,18,19} A. AIN,^{20,21} P. AJITH,²² T. AKUTSU,^{23,24} K. M. ALEMAN,²⁵ G. ALLEN,²⁶ A. ALLOCCA,^{27,5} P. A. ALTIN,⁹ A. AMATO,²⁸ S. ANAND,¹ A. ANANYEVA,¹ S. B. ANDERSON,¹ W. G. ANDERSON,²⁹ M. ANDO,^{30,31,23} S. V. ANGELOVA,³² S. ANSOLDI,^{33,34} J. M. ANTELLIS,³⁵ S. ANTIER,³⁶ S. APPERT,¹ KOYA ARAI,³⁷ KOJI ARAI,¹ Y. ARAI,³⁷ S. ARAKI,³⁸ A. ARAYA,³⁹ M. C. ARAYA,¹ J. S. AREEDA,²⁵ M. ARÈNE,³⁶ N. ARITOMI,³⁰ N. ARNAUD,^{40,41} S. M. ARONSON,⁴² K. G. ARUN,⁴³ H. ASADA,⁴⁴ Y. ASALI,⁴⁵ G. ASHTON,⁶ Y. ASO,^{46,47} S. M. ASTON,⁸ P. ASTONE,⁴⁸ F. AUBIN,⁴⁹ P. AUFMUTH,^{10,11} K. AULTONEAL,³⁵ C. AUSTIN,² S. BABAK,³⁶ F. BADARACCO,^{18,19} M. K. M. BADER,⁵⁰ S. BAE,⁵¹ Y. BAE,⁵² A. M. BAER,⁷ S. BAGNASCO,⁵³ Y. BAI,¹ L. BAIOTTI,⁵⁴ J. BAIRD,³⁶ R. BAJPAI,⁵⁵ M. BALL,⁵⁶ G. 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CHUNG,¹³² G. CIANI,^{72,73} P. CIECIELAG,⁷⁶ M. CIEŚLAR,⁷⁶ M. CIFALDI,^{114,115} A. A. CIOBANU,⁷⁸ R. CIOLFI,^{133,73} F. CIPRIANO,⁸⁸ A. CIRONE,^{108,80} F. CLARA,⁶³ E. N. CLARK,¹³⁴ J. A. CLARK,¹⁰² L. CLARKE,¹³⁵ P. CLEARWATER,¹¹¹ S. CLESSE,¹³⁶ F. CLEVA,⁸⁸ E. COCCIA,^{18,19} P.-F. COHADON,⁹⁶ D. E. COHEN,⁴⁰ L. COHEN,² M. COLLEONI,¹³⁷ C. G. COLLETTE,¹³⁸ M. COLPI,^{60,61} C. M. COMPTON,⁶³ M. CONSTANCIO JR.,¹⁶ L. CONTI,⁷³ S. J. COOPER,¹⁴ P. CORBAN,⁸ T. R. CORBITT,² I. CORDERO-CARRIÓN,¹³⁹ S. COREZZI,^{71,70} K. R. CORLEY,⁴⁵ N. CORNISH,⁷⁴ D. CORRE,⁴⁰ A. CORSI,¹⁴⁰ S. CORTESE,⁴¹ C. A. COSTA,¹⁶ R. COTESTA,¹⁰¹ M. W. COUGHLIN,⁵⁹ S. B. COUGHLIN,^{15,17} J.-P. COULON,⁸⁸ S. T. COUNTRYMAN,⁴⁵ B. COUSINS,¹⁴¹ P. COUVARES,¹ P. B. COVAS,¹³⁷ D. M. COWARD,⁹⁰ M. J. COWART,⁸ D. C. COYNE,¹ R. COYNE,¹⁴² J. D. E. CREIGHTON,²⁹ T. D. CREIGHTON,¹⁴³ A. W. CRISWELL,⁵⁹ M. CROQUETTE,⁹⁶ S. G. CROWDER,¹⁴⁴ J. R. CUDELL,⁵⁸ T. J. CULLEN,² A. CUMMING,⁶⁷ R. CUMMINGS,⁶⁷ E. CUOCO,^{41,145,20} M. CURYŁO,⁹⁸ T. 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- B. EDELMAN,⁵⁶ T. B. EDO,^{1,151} O. EDY,¹⁵⁰ A. EFFLER,⁸ S. EGUCHI,¹²¹ J. EICHHOLZ,⁹ S. S. EIKENBERRY,⁴²
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C. GARCÍA-QUIRÓS,¹³⁷ F. GARUFI,^{27,5} B. GATELEY,⁶³ S. GAUDIO,³⁵ V. GAYATHRI,⁴² G. GE,¹⁶⁹ G. GEMME,⁸⁰ A. GENNAI,²⁰
J. GEORGE,⁸¹ L. GERGELY,¹⁷⁰ P. GEWECKE,¹⁴⁸ S. GHONGE,¹⁰² ABHIRUP. GHOSH,¹⁰¹ ARCHISMAN GHOSH,¹⁷¹
SHAON GHOSH,^{29,158} SHROBANA GHOSH,¹⁷ SOURATH GHOSH,⁴² B. GIACOMAZZO,^{60,61,62} L. GIACOPPO,^{93,48} J. A. GIAIME,^{2,8}
K. D. GIARDINA,⁸ D. R. GIBSON,⁸⁴ C. GIER,³² M. GIESLER,⁸⁹ P. GIRI,^{20,21} F. GISSI,⁷⁷ J. GLANZER,² A. E. GLECKL,²⁵
P. GODWIN,¹⁴¹ E. GOETZ,¹⁷² R. GOETZ,⁴² N. GOHLKE,^{10,11} B. GONCHAROV,⁶ G. GONZÁLEZ,² A. GOPAKUMAR,¹⁷³
M. GOSSELIN,⁴¹ R. GOUATY,⁴⁹ B. GRACE,⁹ A. GRADO,^{174,5} M. GRANATA,¹⁵² V. GRANATA,⁹¹ A. GRANT,⁶⁷ S. GRAS,⁶⁵
P. GRASSIA,¹ C. GRAY,⁶³ R. GRAY,⁶⁷ G. GRECO,⁷⁰ A. C. GREEN,⁴² R. GREEN,¹⁷ A. M. GRETARSSON,³⁵
E. M. GRETARSSON,³⁵ D. GRIFFITH,¹ W. GRIFFITHS,¹⁷ H. L. GRIGGS,¹⁰² G. GRIGNANI,^{71,70} A. GRIMALDI,^{175,176}
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A. R. GUIMARAES,² G. GUIXÉ,⁶⁴ H. K. GULATI,⁷⁵ H.-K. GUO,¹⁶⁴ Y. GUO,⁵⁰ ANCHAL GUPTA,¹ ANURADHA GUPTA,¹⁷⁷
P. GUPTA,^{50,117} E. K. GUSTAFSON,¹ R. GUSTAFSON,¹⁷⁸ F. GUZMAN,¹³⁴ S. HA,¹⁷⁹ L. HAEGEL,³⁶ A. HAGIWARA,^{37,180}
S. HAINO,¹²⁸ O. HALIM,^{181,34} E. D. HALL,⁶⁵ E. Z. HAMILTON,¹⁷ G. HAMMOND,⁶⁷ W.-B. HAN,¹⁸² M. HANEY,¹⁵⁵ J. HANKS,⁶³
C. HANNA,¹⁴¹ M. D. HANNAM,¹⁷ O. A. HANNUKSELA,^{117,50,104} H. HANSEN,⁶³ T. J. HANSEN,³⁵ J. HANSON,⁸ T. HARDER,⁸⁸
T. HARDWICK,² K. HARIS,^{50,117,22} J. HARMS,^{18,19} G. M. HARRY,¹⁸³ I. W. HARRY,¹⁵⁰ D. HARTWIG,¹⁴⁸ K. HASEGAWA,³⁷
B. HASKELL,⁷⁶ R. K. HASSKEW,⁸ C.-J. HASTER,⁶⁵ K. HATTORI,¹⁸⁴ K. HAUGHIAN,⁶⁷ H. HAYAKAWA,¹⁸⁵ K. HAYAMA,¹²¹
F. J. HAYES,⁶⁷ J. HEALY,¹¹⁹ A. HEIDMANN,⁹⁶ M. C. HEINTZE,⁸ J. HEINZE,^{10,11} J. HEINZEL,¹⁸⁶ H. HEITMANN,⁸⁸
F. HELLMAN,¹⁸⁷ P. HELLO,⁴⁰ A. F. HELMLING-CORNELL,⁵⁶ G. HEMMING,⁴¹ M. HENDRY,⁶⁷ I. S. HENG,⁶⁷ E. HENNES,⁵⁰
J. HENNIG,^{10,11} M. H. HENNIG,^{10,11} F. HERNANDEZ VIVANCO,⁶ M. HEURS,^{10,11} S. HILD,^{147,50} P. HILL,³² Y. HIMEMOTO,¹⁸⁸
A. S. HINES,¹³⁴ Y. HIRANUMA,¹⁸⁹ N. HIRATA,²³ E. HIROSE,³⁷ W. C. G. HO,¹⁹⁰ S. HOCHHEIM,^{10,11} D. HOFMAN,¹⁵²
J. N. HOHMANN,¹⁴⁸ A. M. HOLGADO,²⁶ N. A. HOLLAND,⁹ I. J. HOLLOWS,¹⁵¹ Z. J. HOLMES,⁷⁸ K. HOLT,⁸ D. E. HOLZ,¹²⁵
Z. HONG,¹⁹¹ P. HOPKINS,¹⁷ J. HOUGH,⁶⁷ E. J. HOWELL,⁹⁰ C. G. HOY,¹⁷ D. HOYLAND,¹⁴ A. HREIBI,^{10,11} B. HSIEH,³⁷
Y. HSU,¹²⁰ G.-Z. HUANG,¹⁹¹ H.-Y. HUANG,¹²⁸ P. HUANG,¹⁶⁹ Y.-C. HUANG,¹²⁴ Y.-J. HUANG,¹²⁸ Y.-W. HUANG,⁶⁵
M. T. HÜBNER,⁶ A. D. HUDDART,¹³⁵ E. A. HUERTA,²⁶ B. HUGHEY,³⁵ D. C. Y. HUI,¹⁹² V. HUI,⁴⁹ S. HUSA,¹³⁷
S. H. HUTTNER,⁶⁷ R. HUXFORD,¹⁴¹ T. HUYNH-DINH,⁸ S. IDE,¹⁹³ B. IDZKOWSKI,⁹⁸ A. IESS,^{114,115} B. IKENOUE,²⁴ S. IMAM,¹⁹¹
K. INAYOSHI,¹⁹⁴ H. INCHAUSPE,⁴² C. INGRAM,⁷⁸ Y. INOUE,¹²⁶ G. INTINI,^{93,48} K. IOKA,¹⁹⁵ M. ISI,⁶⁵ K. ISLEIF,¹⁴⁸ K. ITO,¹⁹⁶
Y. ITOH,^{197,198} B. R. IYER,²² K. IZUMI,¹⁹⁹ V. JABERIANHAMEDAN,⁹⁰ T. JACQMIN,²⁰⁰ S. J. JADHAV,²⁰⁰ S. P. JADHAV,³
A. L. JAMES,¹⁷ A. Z. JAN,¹¹⁹ K. JANI,¹⁰² K. JANSSENS,²⁰¹ N. N. JANTHALUR,²⁰⁰ P. JARANOWSKI,²⁰² D. JARIWALA,⁴²
R. JAUME,¹³⁷ A. C. JENKINS,¹³² C. JEON,²⁰³ M. JEUNON,⁵⁹ W. JIA,⁶⁵ J. JIANG,⁴² H.-B. JIN,^{204,205} G. R. JOHNS,⁷
A. W. JONES,⁹⁰ D. I. JONES,²⁰⁶ J. D. JONES,⁶³ P. JONES,¹⁴ R. JONES,⁶⁷ R. J. G. JONKER,⁵⁰ L. JU,⁹⁰ K. JUNG,¹⁷⁹
P. JUNG,¹⁸⁵ J. JUNKER,^{10,11} K. KAIHOTSU,¹⁹⁶ T. KAJITA,²⁰⁷ M. KAKIZAKI,¹⁸⁴ C. V. KALAGHATGI,¹⁷ V. KALOGERA,¹⁵
B. KAMAI,¹ M. KAMIIZUMI,¹⁸⁵ N. KANDA,^{197,198} S. KANDHASAMY,³ G. KANG,⁵¹ J. B. KANNER,¹ Y. KAO,¹²⁰
S. J. KAPADIA,²² D. P. KAPASI,⁹ C. KARATHANASIS,²⁰⁸ S. KARKI,⁸² R. KASHYAP,¹⁴¹ M. KASPRZACK,¹ W. KASTAUN,^{10,11}
S. KATSANEVAS,⁴¹ E. KATSAVOUNIDIS,⁶⁵ W. KATZMAN,⁸ T. KAUR,⁹⁰ K. KAWABE,⁶³ K. KAWAGUCHI,³⁷ N. KAWAI,²⁰⁹
T. KAWASAKI,³⁰ F. KÉFÉLIAN,⁸⁸ D. KEITEL,¹³⁷ J. S. KEY,²¹⁰ S. KHADKA,⁶⁹ F. Y. KHALILI,⁸³ I. KHAN,^{18,115} S. KHAN,¹⁷
E. A. KHAZANOV,²¹¹ N. KHETAN,^{18,19} M. KHURSEED,⁸¹ N. KJIBUNCHOO,⁹ C. KIM,^{212,203} J. C. KIM,²¹³ J. KIM,²¹⁴
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SORRENTINO,⁸⁰ N. SORRENTINO,^{21,20} H. SOTANI,²⁶⁹ R. SOULARD,⁸⁸ T. SOURADEEP,^{258,3} E. SOWELL,¹⁴⁰ V. SPAGNUOLO,^{147,50} A. P. SPENCER,⁶⁷ M. SPERA,^{72,73} A. K. SRIVASTAVA,⁷⁵ V. SRIVASTAVA,⁵⁷ K. STAATS,¹⁵ C. STACHIE,⁸⁸ D. A. STEER,³⁶ J. STEINLECHNER,^{147,50} S. STEINLECHNER,^{147,50} D. J. STOPS,¹⁴ M. STOVER,¹⁶⁵ K. A. STRAIN,⁶⁷ L. C. STRANG,¹¹¹ G. STRATTA,^{270,87} A. STRUNK,⁶³ R. STURANI,²⁵² A. L. STUVER,¹⁰³ J. SÜDBECK,¹⁴⁸ S. SUDHAGAR,³ V. SUDHIR,⁶⁵ R. SUGIMOTO,^{271,199} H. G. SUH,²⁹ T. Z. SUMMERSCALES,²⁷² H. SUN,⁹⁰ L. SUN,^{9,1} S. SUNIL,⁷⁵ A. SUR,⁷⁶ J. SURESH,^{31,37} P. J. SUTTON,¹⁷ TAKAMASA SUZUKI,¹⁶⁸ TOSHIKAZU SUZUKI,³⁷ B. L. SWINKELS,⁵⁰ M. J. SZCZEPANČYK,⁴² P. SZEWCHYK,⁹⁸ M. TACCA,⁵⁰ H. TAGOSHI,³⁷ S. C. TAIT,⁶⁷ H. TAKAHASHI,²⁷³ R. TAKAHASHI,²³ A. TAKAMORI,³⁹ S. TAKANO,³⁰ H. TAKEDA,³⁰ M. TAKEDA,¹⁹⁷ C. TALBOT,¹ H. TANAKA,²⁷⁴ KAZUYUKI TANAKA,¹⁹⁷ KENTA TANAKA,²⁷⁴ TAIKI TANAKA,³⁷ TAKAHIRO TANAKA,²⁶¹ A. J. TANASIJCZUK,⁹⁷ S. TANIOKA,^{23,47} D. B. TANNER,⁴² D. TAO,¹ A. TAPIA,²⁵ E. N. TAPIA SAN MARTIN,²³ E. N. TAPIA SAN MARTIN,⁵⁰ J. D. TASSON,¹⁸⁶ S. TELADA,²⁷⁵ R. TENORIO,¹³⁷ L. TERKOWSKI,¹⁴⁸ M. TEST,²⁹ M. P. THIRUGNANASAMBANDAM,³ M. THOMAS,⁸ P. THOMAS,⁶³ J. E. THOMPSON,¹⁷ S. R. THONDAPU,⁸¹ K. A. THORNE,⁸ E. THRANE,⁶ SHUBHANSHU TIWARI,¹⁵⁵ SRISHTI TIWARI,¹⁷³ V. TIWARI,¹⁷ K. TOLAND,⁶⁷ A. E. TOLLEY,¹⁵⁰ T. TOMARU,²³ Y. TOMIGAMI,¹⁹⁷ T. TOMURA,¹⁸⁵ M. TONELLI,^{21,20} A. TORRES-FORNÉ,¹¹⁸ C. I. TORRIE,¹ I. TOSTA E MELO,^{112,113} D. TÖYRÄ,⁹ A. TRAPANANTI,^{236,70} F. TRAVASSO,^{70,236} G. TRAYLOR,⁸ M. C. TRINGALI,⁴¹ A. TRIPATHEE,¹⁷⁸ L. TROIANO,^{276,92} A. TROVATO,³⁶ L. TROZZO,¹⁸⁵ R. J. TRUDEAU,¹ D. S. TSAI,¹²⁰ D. TSAI,¹²⁰ K. W. TSANG,^{50,277,117} T. TSANG,¹⁰⁴ J.-S. TSAO,¹⁹¹ M. TSE,⁶⁵ R. TSO,⁸⁹ K. TSUBONO,³⁰ S. TSUCHIDA,¹⁹⁷ L. TSUKADA,³¹ D. TSUNA,³¹ T. TSUTSUI,³¹ T. TSUZUKI,²⁴ M. TURCONI,⁸⁸ D. TUYENBAYEV,¹²⁸ A. S. UBHI,¹⁴ N. UCHIKATA,³⁷ T. UCHIYAMA,¹⁸⁵ R. P. UDALL,^{102,1} A. UEDA,¹⁸⁰ T. UEHARA,^{278,279} K. UENO,³¹ G. UESHIMA,²⁷³ D. UGOLINI,²⁸⁰ C. S. UNNIKRISSNAN,¹⁷³ F. URAGUCHI,²⁴ A. L. URBAN,² T. USHIBA,³⁷ S. A. USMAN,¹²⁵ A. C. UTINA,^{147,50} H. VAHLBRUCH,^{10,11} G. VAJENTE,¹ A. VAJPEYI,⁶ G. VALDES,² M. VALENTINI,^{175,176} V. VALSAN,²⁹ N. VAN BAKEL,⁵⁰ M. VAN BEUZekom,⁵⁰ J. F. J. VAN DEN BRAND,^{147,99,50} C. VAN DEN BROECK,^{117,50} D. C. VANDER-HYDE,⁵⁷ L. VAN DER SCHAAP,⁵⁰ J. V. VAN HEIJNINGEN,^{90,97} M. H. P. M. VAN PUTTEN,²⁸¹ M. VARDARO,^{229,50} A. F. VARGAS,¹¹¹ V. VARMA,⁸⁹ M. VASÚTH,⁶⁸ A. VECCHIO,¹⁴ G. VEDOVATO,⁷³ J. VEITCH,⁶⁷ P. J. VEITCH,⁷⁸ K. VENKATESWARA,²³¹ J. VENNEBERG,^{10,11} G. VENUGOPALAN,¹ D. VERKINDT,⁴⁹ Y. VERMA,⁸¹ D. VESKE,⁴⁵ F. VETRANO,⁸⁶ A. VICERÉ,^{86,87} A. D. VIETS,²³⁵ V. VILLA-ORTEGA,¹⁴⁹ J.-Y. VINET,⁸⁸ S. VITALE,⁶⁵ T. VO,⁵⁷ H. VOCCA,^{71,70} E. R. G. VON REIS,⁶³ C. VORVICK,⁶³ S. P. VYATCHANIN,⁸³ L. E. WADE,¹⁶⁵ M. WADE,¹⁶⁵ K. J. WAGNER,¹¹⁹ R. C. WALET,⁵⁰ M. WALKER,⁷ G. S. WALLACE,³² L. WALLACE,¹ S. WALSH,²⁹ J. WANG,¹⁶⁹ J. Z. WANG,¹⁷⁸ W. H. WANG,¹⁴³ R. L. WARD,⁹ J. WARNER,⁶³ M. WAS,⁴⁹ T. WASHIMI,²³ N. Y. WASHINGTON,¹ J. WATCHI,¹³⁸ B. WEAVER,⁶³ L. WEI,^{10,11} M. WEINERT,^{10,11} A. J. WEINSTEIN,¹ R. WEISS,⁶⁵ C. M. WELLER,²³¹ F. WELLMANN,^{10,11} L. WEN,⁹⁰ P. WESSELS,^{10,11} J. W. WESTHOUSE,³⁵ K. WETTE,⁹ J. T. WHELAN,¹¹⁹ D. D. WHITE,²⁵ B. F. WHITING,⁴² C. WHITTLE,⁶⁵ D. WILKEN,^{10,11} D. WILLIAMS,⁶⁷ M. J. WILLIAMS,⁶⁷ A. R. WILLIAMSON,¹⁵⁰ J. L. WILLIS,¹ B. WILLKE,^{10,11} D. J. WILSON,¹³⁴ W. WINKLER,^{10,11} C. C. WIPF,¹ T. WLODARCZYK,¹⁰¹ G. WOAN,⁶⁷ J. WOEHLE,^{10,11} J. K. WOFFORD,¹¹⁹ I. C. F. WONG,¹⁰⁴ J. WRANGEL,^{10,11} C. WU,¹²⁴ D. S. WU,^{10,11} H. WU,¹²⁴ S. WU,¹²⁴ D. M. WYSOCKI,^{29,119} L. XIAO,¹ W.-R. XU,¹⁹¹ T. YAMADA,²⁷⁴ H. YAMAMOTO,¹ KAZUHIRO YAMAMOTO,¹⁸⁴ KOHEI YAMAMOTO,²⁷⁴ T. YAMAMOTO,¹⁸⁵ K. YAMASHITA,¹⁸⁴ R. YAMAZAKI,¹⁹³ F. W. YANG,¹⁶⁴ L. YANG,¹⁵⁹ YANG YANG,⁴² YI YANG,²⁸² Z. YANG,⁵⁹ M. J. YAP,⁹ D. W. YEELES,¹⁷ A. B. YELIKAR,¹¹⁹ M. YING,¹²⁰ K. YOKOGAWA,¹⁹⁶ J. YOKOYAMA,^{31,30} T. YOKOZAWA,¹⁸⁵ A. YOON,⁷ T. YOSHIOKA,¹⁹⁶ HANG YU,⁸⁹ HAOCUN YU,⁶⁵ H. YUZURIHARA,³⁷ A. ZADROŻNY,²²⁰ M. ZANOLIN,³⁵ S. ZEIDLER,²⁸³ T. ZELENKOVA,⁴¹ J.-P. ZENDRI,⁷³ M. ZEVIN,¹⁵ M. ZHAN,¹⁶⁹ H. ZHANG,¹⁹¹ J. ZHANG,⁹⁰ L. ZHANG,¹ R. ZHANG,⁴² T. ZHANG,¹⁴ C. ZHAO,⁹⁰ G. ZHAO,¹³⁸ YUE ZHAO,¹⁶⁴ YUHAN ZHAO,²³ Z. ZHOU,¹⁵ X. J. ZHU,⁶ Z.-H. ZHU,¹¹⁰ M. E. ZUCKER,^{1,65} AND J. ZWEIZIG¹
THE LIGO SCIENTIFIC COLLABORATION, THE VIRGO COLLABORATION AND THE KAGRA COLLABORATION
D. ANTONOPOULOU,²⁸⁴ Z. ARZOUMANIAN,²⁸⁵ T. ENOTO,²⁸⁶ C. M. ESPINOZA,²⁸⁷ AND S. GUILLOT^{288,289}

¹LIGO, California Institute of Technology, Pasadena, CA 91125, USA

²Louisiana State University, Baton Rouge, LA 70803, USA

³Inter-University Centre for Astronomy and Astrophysics, Pune 411007, India

⁴Dipartimento di Farmacia, Università di Salerno, I-84084 Fisciano, Salerno, Italy

⁵INFN, Sezione di Napoli, Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy

⁶OzGrav, School of Physics & Astronomy, Monash University, Clayton 3800, Victoria, Australia

⁷Christopher Newport University, Newport News, VA 23606, USA

⁸LIGO Livingston Observatory, Livingston, LA 70754, USA

⁹OzGrav, Australian National University, Canberra, Australian Capital Territory 0200, Australia

¹⁰Max Planck Institute for Gravitational Physics (Albert Einstein Institute), D-30167 Hannover, Germany

¹¹Leibniz Universität Hannover, D-30167 Hannover, Germany

- ¹² *University of Cambridge, Cambridge CB2 1TN, United Kingdom*
- ¹³ *Theoretisch-Physikalisches Institut, Friedrich-Schiller-Universität Jena, D-07743 Jena, Germany*
- ¹⁴ *University of Birmingham, Birmingham B15 2TT, United Kingdom*
- ¹⁵ *Center for Interdisciplinary Exploration & Research in Astrophysics (CIERA), Northwestern University, Evanston, IL 60208, USA*
- ¹⁶ *Instituto Nacional de Pesquisas Espaciais, 12227-010 São José dos Campos, São Paulo, Brazil*
- ¹⁷ *Gravity Exploration Institute, Cardiff University, Cardiff CF24 3AA, United Kingdom*
- ¹⁸ *Gran Sasso Science Institute (GSSI), I-67100 L'Aquila, Italy*
- ¹⁹ *INFN, Laboratori Nazionali del Gran Sasso, I-67100 Assergi, Italy*
- ²⁰ *INFN, Sezione di Pisa, I-56127 Pisa, Italy*
- ²¹ *Università di Pisa, I-56127 Pisa, Italy*
- ²² *International Centre for Theoretical Sciences, Tata Institute of Fundamental Research, Bengaluru 560089, India*
- ²³ *Gravitational Wave Science Project, National Astronomical Observatory of Japan (NAOJ), Mitaka City, Tokyo 181-8588, Japan*
- ²⁴ *Advanced Technology Center, National Astronomical Observatory of Japan (NAOJ), Mitaka City, Tokyo 181-8588, Japan*
- ²⁵ *California State University Fullerton, Fullerton, CA 92831, USA*
- ²⁶ *NCSA, University of Illinois at Urbana-Champaign, Urbana, IL 61801, USA*
- ²⁷ *Università di Napoli “Federico II”, Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*
- ²⁸ *Université de Lyon, Université Claude Bernard Lyon 1, CNRS, Institut Lumière Matière, F-69622 Villeurbanne, France*
- ²⁹ *University of Wisconsin-Milwaukee, Milwaukee, WI 53201, USA*
- ³⁰ *Department of Physics, The University of Tokyo, Bunkyo-ku, Tokyo 113-0033, Japan*
- ³¹ *Research Center for the Early Universe (RESCEU), The University of Tokyo, Bunkyo-ku, Tokyo 113-0033, Japan*
- ³² *SUPA, University of Strathclyde, Glasgow G1 1XQ, United Kingdom*
- ³³ *Dipartimento di Matematica e Informatica, Università di Udine, I-33100 Udine, Italy*
- ³⁴ *INFN, Sezione di Trieste, I-34127 Trieste, Italy*
- ³⁵ *Embry-Riddle Aeronautical University, Prescott, AZ 86301, USA*
- ³⁶ *Université de Paris, CNRS, Astroparticule et Cosmologie, F-75006 Paris, France*
- ³⁷ *Institute for Cosmic Ray Research (ICRR), KAGRA Observatory, The University of Tokyo, Kashiwa City, Chiba 277-8582, Japan*
- ³⁸ *Accelerator Laboratory, High Energy Accelerator Research Organization (KEK), Tsukuba City, Ibaraki 305-0801, Japan*
- ³⁹ *Earthquake Research Institute, The University of Tokyo, Bunkyo-ku, Tokyo 113-0032, Japan*
- ⁴⁰ *Université Paris-Saclay, CNRS/IN2P3, IJCLab, 91405 Orsay, France*
- ⁴¹ *European Gravitational Observatory (EGO), I-56021 Cascina, Pisa, Italy*
- ⁴² *University of Florida, Gainesville, FL 32611, USA*
- ⁴³ *Chennai Mathematical Institute, Chennai 603103, India*
- ⁴⁴ *Department of Mathematics and Physics, Hirosaki University, Hirosaki City, Aomori 036-8561, Japan*
- ⁴⁵ *Columbia University, New York, NY 10027, USA*
- ⁴⁶ *Kamioka Branch, National Astronomical Observatory of Japan (NAOJ), Kamioka-cho, Hida City, Gifu 506-1205, Japan*
- ⁴⁷ *The Graduate University for Advanced Studies (SOKENDAI), Mitaka City, Tokyo 181-8588, Japan*
- ⁴⁸ *INFN, Sezione di Roma, I-00185 Roma, Italy*
- ⁴⁹ *Univ. Grenoble Alpes, Laboratoire d’Annecy de Physique des Particules (LAPP), Université Savoie Mont Blanc, CNRS/IN2P3, F-74941 Annecy, France*
- ⁵⁰ *Nikhef, Science Park 105, 1098 XG Amsterdam, Netherlands*
- ⁵¹ *Korea Institute of Science and Technology Information (KISTI), Yuseong-gu, Daejeon 34141, Korea*
- ⁵² *National Institute for Mathematical Sciences, Daejeon 34047, South Korea*
- ⁵³ *INFN Sezione di Torino, I-10125 Torino, Italy*
- ⁵⁴ *International College, Osaka University, Toyonaka City, Osaka 560-0043, Japan*
- ⁵⁵ *School of High Energy Accelerator Science, The Graduate University for Advanced Studies (SOKENDAI), Tsukuba City, Ibaraki 305-0801, Japan*
- ⁵⁶ *University of Oregon, Eugene, OR 97403, USA*
- ⁵⁷ *Syracuse University, Syracuse, NY 13244, USA*
- ⁵⁸ *Université de Liège, B-4000 Liège, Belgium*
- ⁵⁹ *University of Minnesota, Minneapolis, MN 55455, USA*
- ⁶⁰ *Università degli Studi di Milano-Bicocca, I-20126 Milano, Italy*
- ⁶¹ *INFN, Sezione di Milano-Bicocca, I-20126 Milano, Italy*
- ⁶² *INAF, Osservatorio Astronomico di Brera sede di Merate, I-23807 Merate, Lecco, Italy*
- ⁶³ *LIGO Hanford Observatory, Richland, WA 99352, USA*
- ⁶⁴ *Institut de Ciències del Cosmos, Universitat de Barcelona, C/ Martí i Franquès 1, Barcelona, 08028, Spain*
- ⁶⁵ *LIGO, Massachusetts Institute of Technology, Cambridge, MA 02139, USA*

- ⁶⁶ *Dipartimento di Medicina, “Chirurgia e Odontoiatria Scuola Medica Salernitana”, Università di Salerno, I-84081 Baronissi, Salerno, Italy*
- ⁶⁷ *SUPA, University of Glasgow, Glasgow G12 8QQ, United Kingdom*
- ⁶⁸ *Wigner RCP, RMKI, H-1121 Budapest, Konkoly Thege Miklós út 29-33, Hungary*
- ⁶⁹ *Stanford University, Stanford, CA 94305, USA*
- ⁷⁰ *INFN, Sezione di Perugia, I-06123 Perugia, Italy*
- ⁷¹ *Università di Perugia, I-06123 Perugia, Italy*
- ⁷² *Università di Padova, Dipartimento di Fisica e Astronomia, I-35131 Padova, Italy*
- ⁷³ *INFN, Sezione di Padova, I-35131 Padova, Italy*
- ⁷⁴ *Montana State University, Bozeman, MT 59717, USA*
- ⁷⁵ *Institute for Plasma Research, Bhat, Gandhinagar 382428, India*
- ⁷⁶ *Nicolaus Copernicus Astronomical Center, Polish Academy of Sciences, 00-716, Warsaw, Poland*
- ⁷⁷ *Dipartimento di Ingegneria, Università del Sannio, I-82100 Benevento, Italy*
- ⁷⁸ *OzGrav, University of Adelaide, Adelaide, South Australia 5005, Australia*
- ⁷⁹ *California State University, Los Angeles, 5151 State University Dr, Los Angeles, CA 90032, USA*
- ⁸⁰ *INFN, Sezione di Genova, I-16146 Genova, Italy*
- ⁸¹ *RRCAT, Indore, Madhya Pradesh 452013, India*
- ⁸² *Missouri University of Science and Technology, Rolla, MO 65409, USA*
- ⁸³ *Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia*
- ⁸⁴ *SUPA, University of the West of Scotland, Paisley PA1 2BE, United Kingdom*
- ⁸⁵ *Bar-Ilan University, Ramat Gan, 5290002, Israel*
- ⁸⁶ *Università degli Studi di Urbino “Carlo Bo”, I-61029 Urbino, Italy*
- ⁸⁷ *INFN, Sezione di Firenze, I-50019 Sesto Fiorentino, Firenze, Italy*
- ⁸⁸ *Artemis, Université Côte d’Azur, Observatoire Côte d’Azur, CNRS, F-06304 Nice, France*
- ⁸⁹ *Caltech CaRT, Pasadena, CA 91125, USA*
- ⁹⁰ *OzGrav, University of Western Australia, Crawley, Western Australia 6009, Australia*
- ⁹¹ *Dipartimento di Fisica “E.R. Caianiello”, Università di Salerno, I-84084 Fisciano, Salerno, Italy*
- ⁹² *INFN, Sezione di Napoli, Gruppo Collegato di Salerno, Complesso Universitario di Monte S. Angelo, I-80126 Napoli, Italy*
- ⁹³ *Università di Roma “La Sapienza”, I-00185 Roma, Italy*
- ⁹⁴ *Univ Rennes, CNRS, Institut FOTON - UMR6082, F-3500 Rennes, France*
- ⁹⁵ *Indian Institute of Technology Bombay, Powai, Mumbai 400 076, India*
- ⁹⁶ *Laboratoire Kastler Brossel, Sorbonne Université, CNRS, ENS-Université PSL, Collège de France, F-75005 Paris, France*
- ⁹⁷ *Université catholique de Louvain, B-1348 Louvain-la-Neuve, Belgium*
- ⁹⁸ *Astronomical Observatory Warsaw University, 00-478 Warsaw, Poland*
- ⁹⁹ *VU University Amsterdam, 1081 HV Amsterdam, Netherlands*
- ¹⁰⁰ *University of Maryland, College Park, MD 20742, USA*
- ¹⁰¹ *Max Planck Institute for Gravitational Physics (Albert Einstein Institute), D-14476 Potsdam, Germany*
- ¹⁰² *School of Physics, Georgia Institute of Technology, Atlanta, GA 30332, USA*
- ¹⁰³ *Villanova University, 800 Lancaster Ave, Villanova, PA 19085, USA*
- ¹⁰⁴ *Faculty of Science, Department of Physics, The Chinese University of Hong Kong, Shatin, N.T., Hong Kong*
- ¹⁰⁵ *Stony Brook University, Stony Brook, NY 11794, USA*
- ¹⁰⁶ *Center for Computational Astrophysics, Flatiron Institute, New York, NY 10010, USA*
- ¹⁰⁷ *NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA*
- ¹⁰⁸ *Dipartimento di Fisica, Università degli Studi di Genova, I-16146 Genova, Italy*
- ¹⁰⁹ *Tsinghua University, Beijing 100084, China*
- ¹¹⁰ *Department of Astronomy, Beijing Normal University, Beijing 100875, China*
- ¹¹¹ *OzGrav, University of Melbourne, Parkville, Victoria 3010, Australia*
- ¹¹² *Università degli Studi di Sassari, I-07100 Sassari, Italy*
- ¹¹³ *INFN, Laboratori Nazionali del Sud, I-95125 Catania, Italy*
- ¹¹⁴ *Università di Roma Tor Vergata, I-00133 Roma, Italy*
- ¹¹⁵ *INFN, Sezione di Roma Tor Vergata, I-00133 Roma, Italy*
- ¹¹⁶ *University of Sannio at Benevento, I-82100 Benevento, Italy and INFN, Sezione di Napoli, I-80100 Napoli, Italy*
- ¹¹⁷ *Institute for Gravitational and Subatomic Physics (GRASP), Utrecht University, Princetonplein 1, 3584 CC Utrecht, Netherlands*
- ¹¹⁸ *Departamento de Astronomía y Astrofísica, Universitat de València, E-46100 Burjassot, València, Spain*
- ¹¹⁹ *Rochester Institute of Technology, Rochester, NY 14623, USA*
- ¹²⁰ *National Tsing Hua University, Hsinchu City, 30013 Taiwan, Republic of China*
- ¹²¹ *Department of Applied Physics, Fukuoka University, Jonan, Fukuoka City, Fukuoka 814-0180, Japan*

- ¹²²*OzGrav, Charles Sturt University, Wagga Wagga, New South Wales 2678, Australia*
- ¹²³*Department of Physics, Tamkang University, Danshui Dist., New Taipei City 25137, Taiwan*
- ¹²⁴*Department of Physics and Institute of Astronomy, National Tsing Hua University, Hsinchu 30013, Taiwan*
- ¹²⁵*University of Chicago, Chicago, IL 60637, USA*
- ¹²⁶*Department of Physics, Center for High Energy and High Field Physics, National Central University, Zhongli District, Taoyuan City 32001, Taiwan*
- ¹²⁷*Dipartimento di Ingegneria Industriale (DIIN), Università di Salerno, I-84084 Fisciano, Salerno, Italy*
- ¹²⁸*Institute of Physics, Academia Sinica, Nankang, Taipei 11529, Taiwan*
- ¹²⁹*Institut de Physique des 2 Infinis de Lyon (IP2I), CNRS/IN2P3, Université de Lyon, Université Claude Bernard Lyon 1, F-69622 Villeurbanne, France*
- ¹³⁰*Seoul National University, Seoul 08826, South Korea*
- ¹³¹*Pusan National University, Busan 46241, South Korea*
- ¹³²*King's College London, University of London, London WC2R 2LS, United Kingdom*
- ¹³³*INAF, Osservatorio Astronomico di Padova, I-35122 Padova, Italy*
- ¹³⁴*University of Arizona, Tucson, AZ 85721, USA*
- ¹³⁵*Rutherford Appleton Laboratory, Didcot OX11 0DE, United Kingdom*
- ¹³⁶*Université libre de Bruxelles, Avenue Franklin Roosevelt 50 - 1050 Bruxelles, Belgium*
- ¹³⁷*Universitat de les Illes Balears, IAC3—IEEC, E-07122 Palma de Mallorca, Spain*
- ¹³⁸*Université Libre de Bruxelles, Brussels 1050, Belgium*
- ¹³⁹*Departamento de Matemáticas, Universitat de València, E-46100 Burjassot, València, Spain*
- ¹⁴⁰*Texas Tech University, Lubbock, TX 79409, USA*
- ¹⁴¹*The Pennsylvania State University, University Park, PA 16802, USA*
- ¹⁴²*University of Rhode Island, Kingston, RI 02881, USA*
- ¹⁴³*The University of Texas Rio Grande Valley, Brownsville, TX 78520, USA*
- ¹⁴⁴*Bellevue College, Bellevue, WA 98007, USA*
- ¹⁴⁵*Scuola Normale Superiore, Piazza dei Cavalieri, 7 - 56126 Pisa, Italy*
- ¹⁴⁶*MTA-ELTE Astrophysics Research Group, Institute of Physics, Eötvös University, Budapest 1117, Hungary*
- ¹⁴⁷*Maastricht University, 6200 MD, Maastricht, Netherlands*
- ¹⁴⁸*Universität Hamburg, D-22761 Hamburg, Germany*
- ¹⁴⁹*IGFAE, Campus Sur, Universidade de Santiago de Compostela, 15782 Spain*
- ¹⁵⁰*University of Portsmouth, Portsmouth, PO1 3FX, United Kingdom*
- ¹⁵¹*The University of Sheffield, Sheffield S10 2TN, United Kingdom*
- ¹⁵²*Laboratoire des Matériaux Avancés (LMA), Institut de Physique des 2 Infinis (IP2I) de Lyon, CNRS/IN2P3, Université de Lyon, Université Claude Bernard Lyon 1, F-69622 Villeurbanne, France*
- ¹⁵³*Dipartimento di Scienze Matematiche, Fisiche e Informatiche, Università di Parma, I-43124 Parma, Italy*
- ¹⁵⁴*INFN, Sezione di Milano Bicocca, Gruppo Collegato di Parma, I-43124 Parma, Italy*
- ¹⁵⁵*Physik-Institut, University of Zurich, Winterthurerstrasse 190, 8057 Zurich, Switzerland*
- ¹⁵⁶*Université de Strasbourg, CNRS, IPHC UMR 7178, F-67000 Strasbourg, France*
- ¹⁵⁷*West Virginia University, Morgantown, WV 26506, USA*
- ¹⁵⁸*Montclair State University, Montclair, NJ 07043, USA*
- ¹⁵⁹*Colorado State University, Fort Collins, CO 80523, USA*
- ¹⁶⁰*Institute for Nuclear Research, Hungarian Academy of Sciences, Bem t'er 18/c, H-4026 Debrecen, Hungary*
- ¹⁶¹*CNR-SPIN, c/o Università di Salerno, I-84084 Fisciano, Salerno, Italy*
- ¹⁶²*Scuola di Ingegneria, Università della Basilicata, I-85100 Potenza, Italy*
- ¹⁶³*Osservatori Astronomic, Universitat de València, E-46980 Paterna, València, Spain*
- ¹⁶⁴*The University of Utah, Salt Lake City, UT 84112, USA*
- ¹⁶⁵*Kenyon College, Gambier, OH 43022, USA*
- ¹⁶⁶*Vrije Universiteit Amsterdam, 1081 HV, Amsterdam, Netherlands*
- ¹⁶⁷*Department of Astronomy, The University of Tokyo, Mitaka City, Tokyo 181-8588, Japan*
- ¹⁶⁸*Faculty of Engineering, Niigata University, Nishi-ku, Niigata City, Niigata 950-2181, Japan*
- ¹⁶⁹*State Key Laboratory of Magnetic Resonance and Atomic and Molecular Physics, Innovation Academy for Precision Measurement Science and Technology (APM), Chinese Academy of Sciences, Xiao Hong Shan, Wuhan 430071, China*
- ¹⁷⁰*University of Szeged, Dóm tér 9, Szeged 6720, Hungary*
- ¹⁷¹*Universiteit Gent, B-9000 Gent, Belgium*
- ¹⁷²*University of British Columbia, Vancouver, BC V6T 1Z4, Canada*
- ¹⁷³*Tata Institute of Fundamental Research, Mumbai 400005, India*
- ¹⁷⁴*INAF, Osservatorio Astronomico di Capodimonte, I-80131 Napoli, Italy*

- ¹⁷⁵ *Università di Trento, Dipartimento di Fisica, I-38123 Povo, Trento, Italy*
- ¹⁷⁶ *INFN, Trento Institute for Fundamental Physics and Applications, I-38123 Povo, Trento, Italy*
- ¹⁷⁷ *The University of Mississippi, University, MS 38677, USA*
- ¹⁷⁸ *University of Michigan, Ann Arbor, MI 48109, USA*
- ¹⁷⁹ *Department of Physics, School of Natural Science, Ulsan National Institute of Science and Technology (UNIST), Ulsan-gun, Ulsan 44919, Korea*
- ¹⁸⁰ *Applied Research Laboratory, High Energy Accelerator Research Organization (KEK), Tsukuba City, Ibaraki 305-0801, Japan*
- ¹⁸¹ *Dipartimento di Fisica, Università di Trieste, I-34127 Trieste, Italy*
- ¹⁸² *Chinese Academy of Sciences, Shanghai Astronomical Observatory, Shanghai 200030, China*
- ¹⁸³ *American University, Washington, D.C. 20016, USA*
- ¹⁸⁴ *Faculty of Science, University of Toyama, Toyama City, Toyama 930-8555, Japan*
- ¹⁸⁵ *Institute for Cosmic Ray Research (ICRR), KAGRA Observatory, The University of Tokyo, Kamioka-cho, Gifu City, Gifu 506-1205, Japan*
- ¹⁸⁶ *Carleton College, Northfield, MN 55057, USA*
- ¹⁸⁷ *University of California, Berkeley, CA 94720, USA*
- ¹⁸⁸ *College of Industrial Technology, Nihon University, Narashino City, Chiba 275-8575, Japan*
- ¹⁸⁹ *Graduate School of Science and Technology, Niigata University, Nishi-ku, Niigata City, Niigata 950-2181, Japan*
- ¹⁹⁰ *Department of Physics and Astronomy, Haverford College, 370 Lancaster Ave, Haverford, PA 19041, USA*
- ¹⁹¹ *Department of Physics, National Taiwan Normal University, sec. 4, Taipei 116, Taiwan*
- ¹⁹² *Astronomy & Space Science, Chungnam National University, Yuseong-gu, Daejeon 34134, Korea, Korea*
- ¹⁹³ *Department of Physics and Mathematics, Aoyama Gakuin University, Sagami-hara City, Kanagawa 252-5258, Japan*
- ¹⁹⁴ *Kavli Institute for Astronomy and Astrophysics, Peking University, Haidian District, Beijing 100871, China*
- ¹⁹⁵ *Yukawa Institute for Theoretical Physics (YITP), Kyoto University, Sakyo-ku, Kyoto City, Kyoto 606-8502, Japan*
- ¹⁹⁶ *Graduate School of Science and Engineering, University of Toyama, Toyama City, Toyama 930-8555, Japan*
- ¹⁹⁷ *Department of Physics, Graduate School of Science, Osaka City University, Sumiyoshi-ku, Osaka City, Osaka 558-8585, Japan*
- ¹⁹⁸ *Nambu Yoichiro Institute of Theoretical and Experimental Physics (NITEP), Osaka City University, Sumiyoshi-ku, Osaka City, Osaka 558-8585, Japan*
- ¹⁹⁹ *Institute of Space and Astronautical Science (JAXA), Chuo-ku, Sagami-hara City, Kanagawa 252-0222, Japan*
- ²⁰⁰ *Directorate of Construction, Services & Estate Management, Mumbai 400094 India*
- ²⁰¹ *Universiteit Antwerpen, Prinsstraat 13, 2000 Antwerpen, Belgium*
- ²⁰² *University of Białystok, 15-424 Białystok, Poland*
- ²⁰³ *Department of Physics, Ewha Womans University, Seodaemun-gu, Seoul 03760, Korea*
- ²⁰⁴ *National Astronomical Observatories, Chinese Academic of Sciences, Chaoyang District, Beijing, China*
- ²⁰⁵ *School of Astronomy and Space Science, University of Chinese Academy of Sciences, Chaoyang District, Beijing, China*
- ²⁰⁶ *University of Southampton, Southampton SO17 1BJ, United Kingdom*
- ²⁰⁷ *Institute for Cosmic Ray Research (ICRR), The University of Tokyo, Kashiwa City, Chiba 277-8582, Japan*
- ²⁰⁸ *Institut de Física d'Altes Energies (IFAE), Barcelona Institute of Science and Technology, and ICREA, E-08193 Barcelona, Spain*
- ²⁰⁹ *Graduate School of Science and Technology, Tokyo Institute of Technology, Meguro-ku, Tokyo 152-8551, Japan*
- ²¹⁰ *University of Washington Bothell, Bothell, WA 98011, USA*
- ²¹¹ *Institute of Applied Physics, Nizhny Novgorod, 603950, Russia*
- ²¹² *Ewha Womans University, Seoul 03760, South Korea*
- ²¹³ *Inje University Gimhae, South Gyeongsang 50834, South Korea*
- ²¹⁴ *Department of Physics, Myongji University, Yongin 17058, Korea*
- ²¹⁵ *Korea Astronomy and Space Science Institute (KASI), Yuseong-gu, Daejeon 34055, Korea*
- ²¹⁶ *Department of Physical Science, Hiroshima University, Higashihiroshima City, Hiroshima 730-0213, Japan*
- ²¹⁷ *Bard College, 30 Campus Rd, Annandale-On-Hudson, NY 12504, USA*
- ²¹⁸ *Institute for Cosmic Ray Research (ICRR), Research Center for Cosmic Neutrinos (RCCN), The University of Tokyo, Kamioka-cho, Gifu City, Gifu 506-1205, Japan*
- ²¹⁹ *Institute of Mathematics, Polish Academy of Sciences, 00656 Warsaw, Poland*
- ²²⁰ *National Center for Nuclear Research, 05-400 Świerk-Otwock, Poland*
- ²²¹ *Cornell University, Ithaca, NY 14850, USA*
- ²²² *Institute for Advanced Research, Nagoya University, Furocho, Chikusa-ku, Nagoya City, Aichi 464-8602, Japan*
- ²²³ *Université de Montréal/Polytechnique, Montreal, Quebec H3T 1J4, Canada*
- ²²⁴ *Laboratoire Lagrange, Université Côte d'Azur, Observatoire Côte d'Azur, CNRS, F-06304 Nice, France*
- ²²⁵ *Department of Physics, University of Texas, Austin, TX 78712, USA*
- ²²⁶ *Department of Physics, Hanyang University, Seoul 04763, Korea*
- ²²⁷ *NAVIER, École des Ponts, Univ Gustave Eiffel, CNRS, Marne-la-Vallée, France*

- ²²⁸ *National Center for High-performance computing, National Applied Research Laboratories, Hsinchu Science Park, Hsinchu City 30076, Taiwan*
- ²²⁹ *Institute for High-Energy Physics, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, Netherlands*
- ²³⁰ *NASA Marshall Space Flight Center, Huntsville, AL 35811, USA*
- ²³¹ *University of Washington, Seattle, WA 98195, USA*
- ²³² *Dipartimento di Matematica e Fisica, Università degli Studi Roma Tre, I-00146 Roma, Italy*
- ²³³ *INFN, Sezione di Roma Tre, I-00146 Roma, Italy*
- ²³⁴ *ESPCI, CNRS, F-75005 Paris, France*
- ²³⁵ *Concordia University Wisconsin, Mequon, WI 53097, USA*
- ²³⁶ *Università di Camerino, Dipartimento di Fisica, I-62032 Camerino, Italy*
- ²³⁷ *Southern University and A&M College, Baton Rouge, LA 70813, USA*
- ²³⁸ *Centre Scientifique de Monaco, 8 quai Antoine 1er, MC-98000, Monaco*
- ²³⁹ *Institute for Photon Science and Technology, The University of Tokyo, Bunkyo-ku, Tokyo 113-8656, Japan*
- ²⁴⁰ *Indian Institute of Technology Madras, Chennai 600036, India*
- ²⁴¹ *Saha Institute of Nuclear Physics, Bidhannagar, West Bengal 700064, India*
- ²⁴² *The Applied Electromagnetic Research Institute, National Institute of Information and Communications Technology (NICT), Koganei City, Tokyo 184-8795, Japan*
- ²⁴³ *Institut des Hautes Etudes Scientifiques, F-91440 Bures-sur-Yvette, France*
- ²⁴⁴ *Faculty of Law, Ryukoku University, Fushimi-ku, Kyoto City, Kyoto 612-8577, Japan*
- ²⁴⁵ *Indian Institute of Science Education and Research, Kolkata, Mohanpur, West Bengal 741252, India*
- ²⁴⁶ *Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, Netherlands*
- ²⁴⁷ *Department of Physics, University of Notre Dame, Notre Dame, IN 46556, USA*
- ²⁴⁸ *Department of Physics, National Tsing Hua University, Hsinchu 30013, Taiwan*
- ²⁴⁹ *GRAPPA, Anton Pannekoek Institute for Astronomy and Institute for High-Energy Physics, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, Netherlands*
- ²⁵⁰ *Consiglio Nazionale delle Ricerche - Istituto dei Sistemi Complessi, Piazzale Aldo Moro 5, I-00185 Roma, Italy*
- ²⁵¹ *Hobart and William Smith Colleges, Geneva, NY 14456, USA*
- ²⁵² *International Institute of Physics, Universidade Federal do Rio Grande do Norte, Natal RN 59078-970, Brazil*
- ²⁵³ *Museo Storico della Fisica e Centro Studi e Ricerche “Enrico Fermi”, I-00184 Roma, Italy*
- ²⁵⁴ *Department of Engineering, University of Sannio, Benevento 82100, Italy*
- ²⁵⁵ *Lancaster University, Lancaster LA1 4YW, United Kingdom*
- ²⁵⁶ *OzGrav, Swinburne University of Technology, Hawthorn VIC 3122, Australia*
- ²⁵⁷ *Università di Trento, Dipartimento di Matematica, I-38123 Povo, Trento, Italy*
- ²⁵⁸ *Indian Institute of Science Education and Research, Pune, Maharashtra 411008, India*
- ²⁵⁹ *Dipartimento di Fisica, Università degli Studi di Torino, I-10125 Torino, Italy*
- ²⁶⁰ *Indian Institute of Technology, Palaj, Gandhinagar, Gujarat 382355, India*
- ²⁶¹ *Department of Physics, Kyoto University, Sakyou-ku, Kyoto City, Kyoto 606-8502, Japan*
- ²⁶² *Department of Electronic Control Engineering, National Institute of Technology, Nagaoka College, Nagaoka City, Niigata 940-8532, Japan*
- ²⁶³ *Centro de Astrofísica e Gravitação (CENTRA), Departamento de Física, Instituto Superior Técnico, Universidade de Lisboa, 1049-001 Lisboa, Portugal*
- ²⁶⁴ *Marquette University, 11420 W. Clybourn St., Milwaukee, WI 53233, USA*
- ²⁶⁵ *Graduate School of Science and Engineering, Hosei University, Koganei City, Tokyo 184-8584, Japan*
- ²⁶⁶ *Faculty of Science, Toho University, Funabashi City, Chiba 274-8510, Japan*
- ²⁶⁷ *Faculty of Information Science and Technology, Osaka Institute of Technology, Hirakata City, Osaka 573-0196, Japan*
- ²⁶⁸ *Indian Institute of Technology Hyderabad, Sangareddy, Khandi, Telangana 502285, India*
- ²⁶⁹ *iTHEMS (Interdisciplinary Theoretical and Mathematical Sciences Program), The Institute of Physical and Chemical Research (RIKEN), Wako, Saitama 351-0198, Japan*
- ²⁷⁰ *INAF, Osservatorio di Astrofisica e Scienza dello Spazio, I-40129 Bologna, Italy*
- ²⁷¹ *Department of Space and Astronautical Science, The Graduate University for Advanced Studies (SOKENDAI), Sagami-hara, Kanagawa 252-5210, Japan*
- ²⁷² *Andrews University, Berrien Springs, MI 49104, USA*
- ²⁷³ *Department of Information and Management Systems Engineering, Nagaoka University of Technology, Nagaoka City, Niigata 940-2188, Japan*
- ²⁷⁴ *Institute for Cosmic Ray Research (ICRR), Research Center for Cosmic Neutrinos (RCCN), The University of Tokyo, Kashiwa City, Chiba 277-8582, Japan*
- ²⁷⁵ *National Metrology Institute of Japan, National Institute of Advanced Industrial Science and Technology, Tsukuba City, Ibaraki 305-8568, Japan*

- ²⁷⁶ *Dipartimento di Scienze Aziendali - Management and Innovation Systems (DISA-MIS), Università di Salerno, I-84084 Fisciano, Salerno, Italy*
- ²⁷⁷ *Van Swinderen Institute for Particle Physics and Gravity, University of Groningen, Nijenborgh 4, 9747 AG Groningen, Netherlands*
- ²⁷⁸ *Department of Communications Engineering, National Defense Academy of Japan, Yokosuka City, Kanagawa 239-8686, Japan*
- ²⁷⁹ *Department of Physics, University of Florida, Gainesville, FL 32611, USA*
- ²⁸⁰ *Trinity University, San Antonio, TX 78212, USA*
- ²⁸¹ *Department of Physics and Astronomy, Sejong University, Gwangjin-gu, Seoul 143-747, Korea*
- ²⁸² *Department of Electrophysics, National Chiao Tung University, Hsinchu, Taiwan*
- ²⁸³ *Department of Physics, Rikkyo University, Toshima-ku, Tokyo 171-8501, Japan*
- ²⁸⁴ *Nicolaus Copernicus Astronomical Center, Polish Academy of Sciences, ul. Bartycka 18, 00-716 Warsaw, Poland*
- ²⁸⁵ *X-Ray Astrophysics Laboratory, NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA*
- ²⁸⁶ *Extreme Natural Phenomena RIKEN Hakubi Research Team, RIKEN Cluster for Pioneering Research, 2-1 Hirasawa, Wako, Saitama 351-0198, Japan*
- ²⁸⁷ *Departamento de Física, Universidad de Santiago de Chile, Avenida Ecuador 3493, 9170124 Estación Central, Santiago, Chile*
- ²⁸⁸ *IRAP, CNRS, 9 avenue du Colonel Roche, BP 44346, F-31028 Toulouse Cedex 4, France*
- ²⁸⁹ *Université de Toulouse, CNES, UPS-OMP, F-31028 Toulouse, France*

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ABSTRACT

We present a search for continuous gravitational-wave signals from the young, energetic X-ray pulsar PSR J0537–6910 using data from the second and third observing runs of LIGO and Virgo. The search is enabled by a contemporaneous timing ephemeris obtained using *NICER* data. The *NICER* ephemeris has also been extended through 2020 October and includes three new glitches. PSR J0537–6910 has the largest spin-down luminosity of any pulsar and is highly active with regards to glitches. Analyses of its long-term and inter-glitch braking indices provided intriguing evidence that its spin-down energy budget may include gravitational-wave emission from a time-varying mass quadrupole moment. Its 62 Hz rotation frequency also puts its possible gravitational-wave emission in the most sensitive band of LIGO/Virgo detectors. Motivated by these considerations, we search for gravitational-wave emission at both once and twice the rotation frequency. We find no signal, however, and report our upper limits. Assuming a rigidly rotating triaxial star, our constraints reach below the gravitational-wave spin-down limit for this star for the first time by more than a factor of two and limit gravitational waves from the $l = m = 2$ mode to account for less than 14% of the spin-down energy budget. The fiducial equatorial ellipticity is limited to less than about 3×10^{-5} , which is the third best constraint for any young pulsar.

Keywords: gravitational waves — pulsars: general — pulsars: individual (PSR J0537–6910) — stars: neutron

1. INTRODUCTION

The young (1–5 kyr) energetic pulsar PSR J0537–6910 (Wang & Gotthelf 1998; Chen et al. 2006) resides in the Large Magellanic Cloud at a distance of 49.6 kpc (Pietrzyński et al. 2019). Its pulsations are only detectable at X-ray energies, and the pulsar was first observed by Marshall et al. (1998) using the *Rossi X-ray Timing Explorer* (*RXTE*) during searches for pulsations from the remnant of SN1987A. Further observations with *RXTE*, prior to its decommissioning in early 2012, revealed that PSR J0537–6910 often undergoes sudden changes in rotation frequency, i.e., glitches, at a rate

of more than three per year, and exhibits interesting inter-glitch behavior (Marshall et al. 2004; Middleditch et al. 2006; Andersson et al. 2018; Antonopoulou et al. 2018; Ferdman et al. 2018). Observations of the pulsar resumed in 2017 using the *Neutron star Interior Composition Explorer* (*NICER*) on board the International Space Station (Gendreau et al. 2012), and these observations from 2017–2020 revealed further glitches and continuation of timing behavior seen with *RXTE* (Ho et al. 2020b).

PSR J0537–6910 is a particularly intriguing potential gravitational-wave source. It is the fastest-spinning known young pulsar (with rotation frequency $f_{\text{rot}} = 62$ Hz), which places its gravitational-wave frequency f (e.g., at twice f_{rot} ; see Section 2.1) in the most sensitive band of ground-based gravitational-wave detec-

* Deceased, August 2020.

tors. PSR J0537–6910 also has the highest spin-down luminosity ($\dot{E} = 4.9 \times 10^{38} \text{ erg s}^{-1}$) among the ~ 2900 known pulsars in the ATNF Pulsar Catalogue (Manchester et al. 2005). Its spin-down behavior appears to be driven by a process other than pure electromagnetic dipole radiation loss (at constant stellar magnetic field and moment of inertia). Specifically, its (long-term) braking index $n \equiv f_{\text{rot}} \ddot{f}_{\text{rot}} / \dot{f}_{\text{rot}}^2 = -1.25 \pm 0.01$, as measured over more than 20 yr (Ho et al. 2020b), indicates an accelerating spin-down rate and significantly deviates from the canonical value of 3 for dipole radiation (Shapiro & Teukolsky 1983), as is true for most pulsars with measurable braking indices.

More importantly, observations of PSR J0537–6910 show the pulsar’s (short-term) interglitch braking index n_{ig} , as measured during intervals between ~ 50 glitches, has values typically > 10 and approaches an asymptotic value of $\lesssim 7$ at long times after a glitch, i.e., when the effects of a preceding glitch are diminished (see Figure 1; see also Andersson et al. 2018). It is this behavior that provides tantalizing suggestions that perhaps PSR J0537–6910 is losing some of its rotational energy to gravitational-wave emission. In particular, a slightly deformed pulsar can emit gravitational waves that results in $n = 5$, and a r-mode fluid oscillation in a pulsar can emit gravitational waves that results in $n = 7$ (see, e.g., Riles 2017; Andersson et al. 2018; Glampedakis & Gualtieri 2018).

In this work, we search for mass quadrupolar gravitational-wave emission from PSR J0537–6910 that follows the same phase as that of the pulsar’s rotation. Previously, data from initial LIGO’s fifth and sixth science runs (S5 and S6) and Virgo’s second and fourth science runs (VSR2 and VSR4), in conjunction with *RXTE* timing measurements, were used to set limits on gravitational-wave emission by PSR J0537–6910 that closely approached the spin-down limit (Abbott et al. 2010; Aasi et al. 2014). Here, we analyze data from the second and third observing runs (O2 and O3) of LIGO and Virgo, tracking the rotation phase with the contemporaneous *NICER* timing ephemeris. In doing so, we also provide an updated ephemeris that includes the latest six months of *NICER* observations of PSR J0537–6910. Investigations of r-mode gravitational-wave emission ($n = 7$) are not presented here; such searches are more technically challenging and require different methods that search over a range of frequencies (see, e.g., Mytidis et al. 2015, 2019; Abbott et al. 2019b; Fesik & Papa 2020a,b) due to uncertainty in gravitational-wave frequency for a given rotation frequency (Andersson et al. 2014; Idrisy et al. 2015; Caride et al. 2019). Nevertheless, we are able to reach below

the spin-down limit of PSR J0537–6910 for the first time, which means that the minimum amplitude we could detect in our analysis is lower than the one given by assuming all of the pulsar’s rotational energy loss is converted to gravitational waves (see Section 2.1). In other words, we can now obtain physically meaningful constraints.

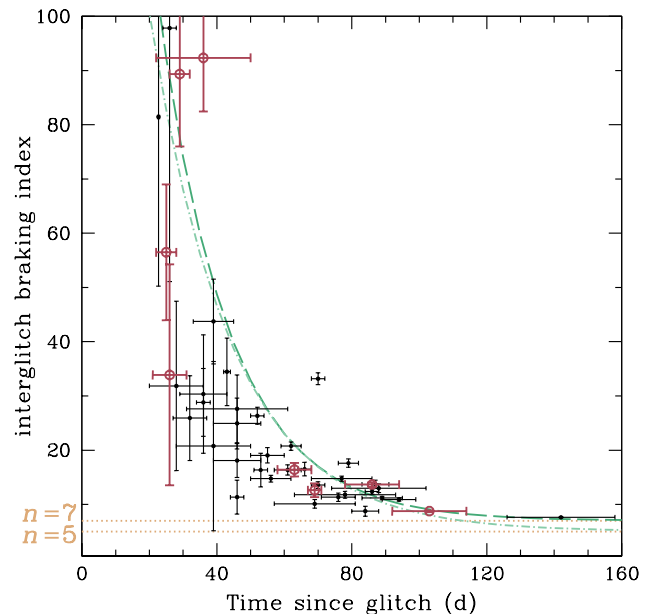


Figure 1. Interglitch braking index n_{ig} calculated from the spin parameters of each segment between glitches as a function of time since the last glitch. Large and small circles denote *NICER* and *RXTE* values, respectively, with the former from Tables 1 and 2 and from Ho et al. (2020b) and latter from Antonopoulou et al. (2018). Errors in n_{ig} are 1σ uncertainty. Orange horizontal dotted lines indicate braking index $n = 5$ and 7, which are expected for pulsar spin-down by gravitational-wave emission due to an ellipticity and r-mode oscillation, respectively. Green dot-dashed and dashed lines indicate exponential decay to $n = 5$ with best-fit time-scale of 24 d and to $n = 7$ with best-fit time-scale of 21 d, respectively.

2. SEARCH METHOD

2.1. Model of gravitational-wave emission

The first model considered here allows for gravitational-wave emission at once and twice the spin frequency simultaneously, which has been searched for previously (Pitkin et al. 2015; Abbott et al. 2017, 2019a, 2020), and can result from a triaxial star spinning about an axis that is not its principal axis (Jones 2010, 2015). The amplitudes of each harmonic at once and twice the spin frequency of the star, denoted $h_{21}(t)$ and $h_{22}(t)$,

respectively, can be written as

$$h_{21} = -\frac{C_{21}}{2} \left\{ F_+^D(\alpha, \delta, \psi; t) \sin \iota \cos \iota \cos [\Phi(t) + \Phi_{21}^C] \right. \\ \left. + F_\times^D(\alpha, \delta, \psi; t) \sin \iota \sin [\Phi(t) + \Phi_{21}^C] \right\}, \quad (1)$$

$$h_{22} = -C_{22} \left\{ F_+^D(\alpha, \delta, \psi; t) (1 + \cos^2 \iota) \cos [2\Phi(t) + \Phi_{22}^C] \right. \\ \left. + 2F_\times^D(\alpha, \delta, \psi; t) \cos \iota \sin [2\Phi(t) + \Phi_{22}^C] \right\}. \quad (2)$$

Here, C_{21} and C_{22} are dimensionless constant component amplitudes, and Φ_{21}^C and Φ_{22}^C are phase angles. F_+^D and F_\times^D are antenna or beam functions and describe how the two polarization components of the signal project onto the detector (see, e.g., Jaranowski et al. 1998). Angles (α, δ) are the right ascension and declination of the source, while angles (ι, ψ) specify the orientation of the star's spin axis relative to the observer. $\Phi(t)$ is the rotational phase of the source.

The second model is a special case of the first model and is used for gravitational-wave emission at only twice the rotational frequency ($C_{21} = 0$), implying a triaxial star that is spinning about a principal axis, such as its z-axis. In this case, it is simpler to write the gravitational-wave amplitude in terms of the dimensionless value h_0 , where in equation (2) the substitution $C_{22} = -h_0/2$ would be made (Abbott et al. 2019a) with the sign change just to maintain consistency with the model from Jaranowski et al. (1998). The cause of such gravitational-wave emission is a deviation from axial symmetry, which can be written in terms of a dimensionless equatorial ellipticity ε , defined in terms of the star's principal moments of inertia (I_{xx}, I_{yy}, I_{zz}):

$$\varepsilon \equiv \frac{|I_{xx} - I_{yy}|}{I_{zz}}. \quad (3)$$

The gravitational-wave amplitude is directly proportional to the ellipticity:

$$h_0 = \frac{16\pi^2 G}{c^4} \frac{I_{zz} \varepsilon f_{\text{rot}}^2}{d}, \quad (4)$$

where d is the star's distance from the Earth. When setting upper limits, we use a fiducial value for the z-component of the moment of inertia, i.e., $I_{zz}^{\text{fid}} = 10^{38} \text{ kg m}^2$. The combination of the ellipticity and fiducial moment of inertia can be cast in terms of the mass quadrupole moment of the $l = m = 2$ mode of the star via $Q_{22} = \sqrt{15/8\pi} I_{zz} \varepsilon$ (Owen 2005). The gravitational-wave amplitude h_0 can be compared to the spin-down limit amplitude h_0^{sd} , which is the gravitational-wave amplitude produced assuming that the entire rotational en-

ergy loss of the pulsar is converted into gravitational waves:

$$h_0^{\text{sd}} = \frac{1}{d} \left(\frac{5GI_{zz}}{2c^3} \frac{|\dot{f}_{\text{rot}}|}{f_{\text{rot}}} \right)^{1/2}. \quad (5)$$

Our results for the single harmonic case are quoted in terms of h_0^{sd} .

NICER observations of PSR J0537–6910 allow for the ephemeris of the pulsar to be determined, which means we know the expected signal frequency and its evolution. With this information, we can perform a targeted search for gravitational waves from this pulsar based on the two signal models discussed, with the phase tracking that of the pulsar rotation.

2.2. *NICER* data

In Ho et al. (2020b), timing analysis is performed on *NICER* data of PSR J0537–6910 from 2017 August 17 to 2020 April 25, with eight glitches detected during this timespan and the last three glitches during O3. Here we present an update and results on timing analysis since the work of Ho et al. (2020b). In particular, data from 2020 May 12 to October 29 is analyzed using the methodology as described in Ho et al. (2020b). Our analysis reveals continuing accelerated spin-down (see Table 1) and three subsequent glitches (see Table 2 and Figure 2), including the smallest glitch of PSR J0537–6910 yet detected using *NICER*. Note that the timing model of segment 8 uses three additional subsequent times-of-arrival (TOAs) beyond those in Table 1 of Ho et al. (2020b) and, as a result, the epoch and other parameters of the model differ; e.g., the longer timespan and lower n_{ig} of segment 8 result in a different position in Figure 1 compared to Figure 6 of Ho et al. (2020b). Meanwhile, the relatively short timespan of segment 9 means the timing model for this segment is not able to constrain \dot{f}_{rot} . For the most recent glitch 11, its magnitude is large ($\Delta f_{\text{rot}} = 33.9 \mu\text{Hz}$), which suggests the time to the next glitch will be long ($\sim 200 \pm 20 \text{ d}$; Ho et al. 2020b). If the interglitch period is indeed long, then *NICER* measurements could eventually yield $n_{\text{ig}} \lesssim 7$ for segment 11, which would lend further support for gravitational-wave emission (see Section 1 and Figure 1).

The gravitational-wave search performed here uses the timing model of Ho et al. (2020b). The differences between the model of Ho et al. (2020b) and the model presented here are well within the former's uncertainties, and thus use of the latter would not yield significantly different results.

Table 1. Timing model parameters for segments between epochs of new glitches of PSR J0537–6910. Columns from left to right are segment number, timing model epoch, segment start and end dates, number of times-of-arrival, rotation frequency and its first two time derivatives, interglitch braking index, and timing model residual and goodness-of-fit measure. Number in parentheses is 1σ uncertainty in last digit. Segments 1–7 are presented in [Ho et al. \(2020b\)](#).

Segment	Epoch	Start	End	TOAs	f_{rot}	\dot{f}_{rot}	\ddot{f}_{rot}	n_{ig}	Residual RMS	χ^2/dof
	(MJD)	(MJD)	(MJD)		(Hz)	(10^{-10} Hz s $^{-1}$)	(10^{-20} Hz s $^{-2}$)		(μs)	
8	58931	58871.5	58991.2	17	61.908808739(3)	−1.997535(7)	1.06(8)	16(1)	173.7	9.9
9	59020	58995.6	59046.3	11	61.907273376(2)	−1.99699(4)	[1] ^a	—	147.8	6.7
10	59074	59050.4	59098.7	10	61.906349948(5)	−1.99762(2)	3.6(8)	56(13)	60.9	1.5
11	59129	59108.7	59150.7	11	61.905434556(6)	−1.99809(3)	2.2(13)	34(20)	72.3	2.1

^a \ddot{f}_{rot} is fixed at 10^{-20} Hz s $^{-2}$.

Table 2. Parameters of new glitches of PSR J0537–6910. Columns from left to right are glitch number and epoch, change in rotation phase and changes in rotation frequency and its first two time derivatives at each glitch. Number in parentheses is 1σ uncertainty in last digit. Glitches 1–7 are presented in [Ho et al. \(2020b\)](#).

Glitch	Glitch epoch	$\Delta\phi$	Δf_{rot}	$\Delta\dot{f}_{\text{rot}}$	$\Delta\ddot{f}_{\text{rot}}$
	(MJD)	(cycle)	(μHz)	(10^{-13} Hz s $^{-1}$)	(10^{-20} Hz s $^{-2}$)
8	58868(5)	0.08(12)	24.0(1)	−2.3(6)	−5(1)
9	58993(3)	0.06(12)	0.4(1)	−0.3(8)	—
10	59049(3)	−0.22(2)	8.46(3)	−1.3(5)	—
11	59103(5)	0.42(2)	33.958(7)	−2.0(3)	—

2.3. LIGO and Virgo data

We use a combination of data from the second and third observing runs of the Advanced LIGO ([Aasi et al. 2015](#)) and Virgo ([Acernese et al. 2015](#)) gravitational wave detectors. During O2, LIGO Livingston (L1) and LIGO Hanford (H1) took data from 2016 November 30 to 2017 August 25 and had duty factors of $\sim 57\%$ and $\sim 59\%$, respectively (including commissioning breaks), while Virgo took data from 2017 August 1 to 2017 August 25 with a duty factor of $\sim 85\%$. As noted in Section 2.2, *NICER* data start on 2017 August 17, and thus one set of searches we undertake uses only about six days of O2 data overlapping with the *NICER* data in addition to the O3 data. Alternatively, we can consider a more optimistic and much longer time-series of O2 data by taking advantage of the correlation between glitch size and time-to-next-glitch seen for PSR J0537–6910 ([Middleditch et al. 2006](#); [Antonopoulou et al. 2018](#); [Ferdman et al. 2018](#); [Ho et al. 2020b](#)). Assuming a (unobserved) glitch occurred on 2017 March 22 with the same size as the largest *NICER* glitch (i.e., glitch 2 with $\Delta f_{\text{rot}} = 36 \mu\text{Hz}$), we would expect a subsequent glitch

224 d later (at 68% confidence) on 2017 November 1, which is the earliest estimated date at which glitch 1 occurred (see Figure 2 and [Ho et al. 2020b](#)). Thus 2017 March 22 to November 1 is the longest period over which we would expect PSR J0537–6910 to not have undergone a glitch and the *NICER* ephemeris to be valid. O3 lasted from 2019 April 1 to 2020 March 27, with a one-month pause in data collection in October 2019. The three detectors’ datasets H1, L1, and V1 had duty factors of $\sim 76\%$, $\sim 77\%$, and $\sim 76\%$ respectively during O3.

In the case of a detection, calibration uncertainties limit our ability to provide robust estimates of the amplitude of the gravitational-wave signal and corresponding ellipticity ([Abbott et al. 2017](#)). Even without a detection, these uncertainties affect the estimated instrument sensitivity and inferred upper limits. The uncertainties vary over the course of a run but do not change by large values, so we do not explicitly consider time-dependent calibration uncertainties in our analysis. For further information on O2 calibration techniques, see discussions in [Abbott et al. \(2019a\)](#).

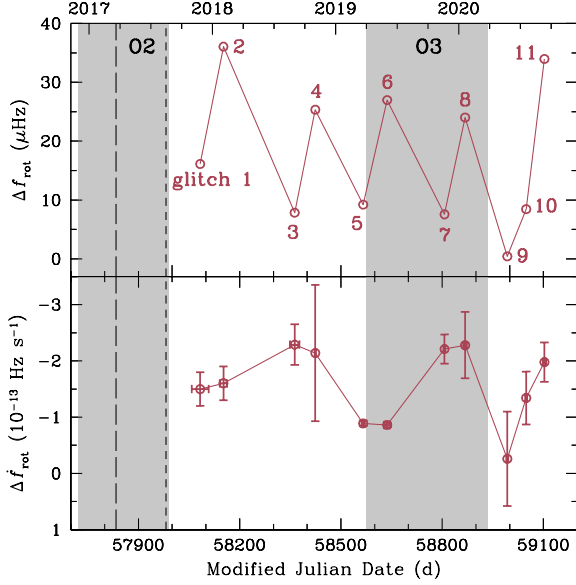


Figure 2. Glitch Δf_{rot} (top) and $\Delta \dot{f}_{\text{rot}}$ (bottom) as functions of time. Glitch numbers and values from Table 2 and Ho et al. (2020b). Errors in $\Delta \dot{f}_{\text{rot}}$ are 1σ uncertainty, while errors in Δf_{rot} are not shown because they are generally smaller than the symbols. Shaded regions denote second observing run (O2) and third observing run (O3) of LIGO/Virgo. Vertical long and short-dashed lines indicate two possible start dates of O2 data used in present work (see Section 2.3).

The full raw strain data from the O2 run is publicly available from the Gravitational Wave Open Science Center¹ (Vallisneri et al. 2015; Abbott et al. 2019c). For the LIGO O3 data set, the analysis uses the “C01” calibration. The C01 calibration has estimated maximum amplitude and phase uncertainties of $\sim 7\%$ and $\sim 4\text{deg}$, respectively (Sun et al. 2020), which we use as conservative estimates of the true calibration uncertainty near the frequencies analyzed here. For the Virgo O3 data set, we use the “V0” calibration with estimated maximum amplitude and phase uncertainties of 5% and 2deg , respectively. We note that the signal frequencies analyzed in L1 and H1 data are close to the US power-line frequencies of 60 and 120 Hz. However, these dis-

turbances do not affect our analysis and results since we consider only a narrow band around the expected signal frequency, as can be seen in the relatively clean amplitude spectral densities of Figure 3.

2.4. Search pipeline

The time-domain Bayesian method performs a coherent analysis of the interferometers’ data, meaning that we analyze the entire data set with an effective single Fourier Transform, thereby preserving the phase information. First, the raw strain data are heterodyned (Dupuis & Woan 2005) using the expected signal phase evolution, known precisely from the electromagnetic timing ephemeris. Then a low-pass filter with a knee frequency of 0.25 Hz is applied, and the data are downsampled so that the sampling time is one minute, compared to 60 microseconds originally. This heterodyning is performed for an expected signal whose frequency is at once or twice the rotational frequency of the pulsar. The heterodyned data is the input to a nested sampling algorithm that is a part of the LALINFERENCE package (Veitch & Vecchio 2010; Veitch et al. 2015), which infers the unknown signal parameters depending on the model of gravitational-wave emission.

PSR J0537–6910 glitched three times over the course of the gravitational-wave observations (see Figure 2). For each glitch, we assume an unknown phase offset between the electromagnetic and gravitational-wave phase. The individual phase offsets of multiple glitches that occurred between O2 and O3 cannot be disentangled, so only one phase offset is included for these glitches. This means that we introduce four additional phase parameters when performing parameter estimation.

We also make use of restricted and unrestricted priors when performing the analysis. In the first case, we use estimates of the orientation of the pulsar relative to the Earth based on a model fit of the observed pulsar wind nebulae torus (Ng & Romani 2008), which imply narrow priors in our analysis on the polarization and inclination angles. From these we use a Gaussian prior on ψ of $2.2864 \pm 0.0384\text{rad}$ and a bimodal Gaussian prior on ι with modes at 1.522 ± 0.016 and $1.620 \pm 0.016\text{rad}$ (see Jones 2015, for reasons behind the bimodality). This range of ι would suggest the pulsar’s rotation axis is almost perpendicular to the line-of-sight, which would in turn lead to a linearly polarized gravitational-wave signal dominated by the ‘+’ polarization mode. The second case assumes a uniform isotropic prior on the axis direction, which therefore does not rely on the above modeling of observations. The initial signal phase and glitch phase offsets all use uniform priors over their full ranges. For the single harmonic search, we parameterize

¹ <https://www.gw-openscience.org/data>

the signals using the mass quadrupole Q_{22} and distance. As a conservative approach, we use an unphysical flat prior on Q_{22} with a lower bound at zero and an upper bound of $5 \times 10^{37} \text{ kg m}^2$, which is well above the largest upper limits found in Abbott et al. (2019a). For the distance, we use a Gaussian prior with mean of 49.59 kpc and standard deviation of 0.55 kpc based on the value given in Pietrzyński et al. (2019), combining the statistical and systematic errors in quadrature. For the dual harmonic search, which uses the amplitudes C_{21} and C_{22} rather than the physical parameters of Q_{22} and d , we use flat priors that are bounded between zero and 1×10^{-22} , which is again well above the limit implied in Abbott et al. (2019a). To analyze multiple detectors’ data sets simultaneously, we combine the product of the likelihoods calculated for each detector (Dupuis & Woan 2005).

The outputs of the analysis are posterior distributions of the parameters of interest, which are $h_0/Q_{22}/\varepsilon$ for the single harmonic search and C_{21} and C_{22} for the dual harmonic search, and of the angles $\cos \iota$ and ψ for both choices of priors. In Section 3, we present results on the amplitude parameters marginalized over the rest of the parameter space.

3. RESULTS

Results from our searches do not show evidence for gravitational-wave emission from PSR J0537–6910 via the two models that we assume. An amplitude spectral density obtained after the heterodyne correction is displayed in Figure 3 for each of the three detectors. If a loud continuous gravitational-wave signal was present, we would expect to see a narrow line feature in the spectrum. The amplitude spectral densities also give an estimation of the sensitivity of the search.

Though no gravitational waves are detected, we can still determine upper limits on possible gravitational-wave emission from PSR J0537–6910. Here we use 95% credible upper bounds on the amplitude parameters based on their marginalized probability distributions. The dimensionless gravitational-wave amplitude h_0 and coefficients C_{21} and C_{22} are constrained for the single and dual harmonic searches, respectively. For the single harmonic search, h_0 can be mapped to a limit on the maximum ellipticity ε using equation (4). In Table 3 we show the different constraints for both searches using all O3 data and the last ~ 6 days of O2 data (see Section 2.3). In addition to the detector calibration uncertainties discussed in Section 2.3, we estimate that the statistical uncertainty on the upper limits due to the use of a finite number of posterior samples is on the order of 1%.

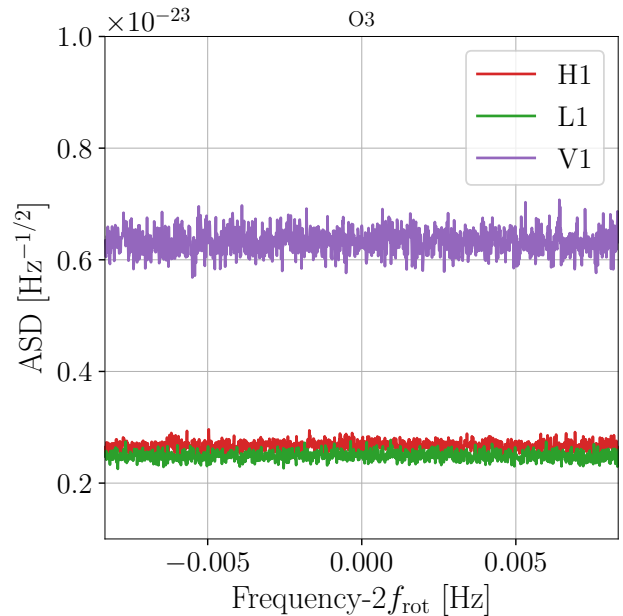


Figure 3. Two-sided amplitude spectral density (ASD) after heterodyning, low pass filtering, and downsampling the raw strain data for the $l = m = 2$ gravitational-wave mode. Different color lines indicate the Hanford (H1), Livingston (L1), and Virgo (V1) detectors.

Table 3. 95% upper limits on gravitational-wave strain, ellipticity, and other quantities based on unrestricted (UR) and restricted (R) choices for priors on polarization and inclination angles. Results here come from analyzing all O3 data and the last 6 days of O2 data.

Prior	$h_0^{95\%}$ (10^{-26})	$\varepsilon^{95\%}$ (10^{-5})	$h_0^{95\%}/h_0^{\text{sd}}$	$C_{21}^{95\%}$ (10^{-26})	$C_{22}^{95\%}$ (10^{-27})
UR	1.1	3.4	0.37	2.2	5.6
R	1.0	3.1	0.33	1.8	5.0

Figure 4 shows the marginalized posterior probability distributions on the pulsar ellipticity and h_0 for the single harmonic search with unrestricted and restricted source orientation priors. The posteriors show significant support at ellipticities of zero, indicating no evidence of a signal at current sensitivities. We therefore show 95% credible upper limits on the ellipticity for both prior choices along with the fiducial spin-down limit.

Figure 5 shows a similar posterior distribution on the dimensionless amplitudes C_{21} and C_{22} for the dual harmonic model. For this model, no evidence of gravitational waves is found, so an upper limit at 95% is indicated in both panels of this figure. The model given

by Equation (1) means that the value of C_{21} becomes completely unconstrained when $\sin \iota = 0$. For the unrestricted orientation prior result shown in the left panel of Figure 5, this leads to a long high amplitude tail in the C_{21} posterior distribution. In Figures 4 and 5, we see that the amplitude posteriors can peak away from zero. This behavior was unsurprising and can occur even for pure Gaussian noise. Even with these peaks, the posteriors are still entirely consistent with zero ellipticity. For example, for the unrestricted posterior distribution shown in Figure 4, a value of zero ellipticity is within the minimum 66% credible interval around the mode.

The results presented above use all O3 data in combination with about 1 week of O2 data, when *NICER* was operating and monitoring PSR J0537–6910. We also conducted searches using only O3 data or using O3 data plus O2 data from 2017 March 22 to the end of O2. The latter analysis assumes no glitches occurred during the additional time and represents the estimated maximum time that can be safely included without a contemporaneous timing model (see Section 2.3). For only O3 data, we obtain h_0 and ε limits that are worse by $\sim 7\%$ for UR and unchanged for R from those shown in Table 3 for which a small amount of O2 data is used. For O3 data plus the extra O2 data, we obtain amplitude limits that are improved by $\lesssim 20\%$ compared to those shown in Table 3.

4. CONCLUSIONS

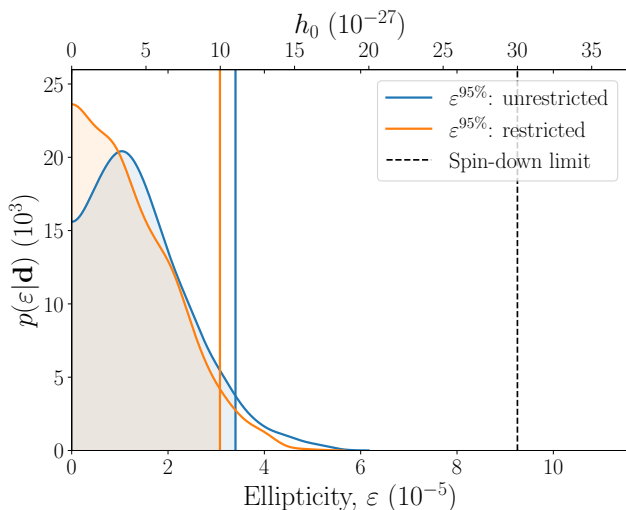


Figure 4. Posterior probability distribution for ellipticity and h_0 for the analyses with unrestricted and restricted priors on the pulsar orientation. The 95% credible upper limits are shown as vertical colored lines, while the spin-down limit is given by the vertical dashed black line.

Using data from LIGO/Virgo’s second and third observing runs, we searched for mass quadrupolar-sourced gravitational waves from the young, dynamic PSR J0537–6910 at once or twice the pulsar’s rotational frequency of 62 Hz. For the first time we reached below the gravitational-wave spin-down limit for PSR J0537–6910 and showed that gravitational-wave emission for a pure $l = m = 2$ mode accounts for less than 14% of the pulsar’s spin-down energy budget. We placed the third most stringent constraint on the ellipticity ($\varepsilon < 3 \times 10^{-5}$) of any young pulsar (behind only the Crab pulsar and B1951+32/J1952+3252; Abbott et al. 2019a, 2020). While this limit is much higher than those of old recycled millisecond pulsars (for which $\varepsilon < 10^{-8}$; Abbott et al. 2020), young pulsars such as PSR J0537–6910 and the Crab pulsar are important because they have much stronger magnetic fields (and are hotter) and thus might have greater ellipticities. The ellipticity constraint of PSR J0537–6910 is also near estimates of the maximum ellipticity that can be sustained by an elastically deformed neutron star crust (Johnson-McDaniel & Owen 2013; Caplan et al. 2018).

PSR J0537–6910 is a frequently glitching pulsar and potential source of continuous gravitational waves. The X-ray data from *NICER* gives us the necessary tools to account for the phase evolution of a gravitational-wave signal over time, which allows us to perform a fully coherent and sensitive search for such a signal. While our multi-messenger analysis focuses on gravitational waves from a time-varying mass quadrupole ($n = 5$), another search could be performed for gravitational waves from a r-mode fluid oscillation ($n = 7$) using wider-band techniques (e.g., Fesik & Papa 2020a,b, using O2 data). The strain sensitivity achieved in our analysis (1×10^{-26}) is also comparable to the $(2 - 3) \times 10^{-26}$ estimated in Andersson et al. (2018) for r-mode emission from PSR J0537–6910.

Finally, from the observed correlation between glitch size and time-to-next-glitch for PSR J0537–6910 (Middleditch et al. 2006; Antonopoulou et al. 2018; Ferdman et al. 2018; Ho et al. 2020b), we can hope to measure in the future low braking indices (7 or even lower) after the largest glitches. As noted above, braking indices of 5 and 7 are predicted by gravitational wave-emitting mechanisms. The observed evolution of n_{ig} to lower values than shown in Figure 1, which may occur after the effects of glitches on the pulsar’s spin-down behavior have decayed, may indicate that gravitational waves are continuously emitted between glitches. On the other hand, glitches may trigger detectable transient gravitational waves (Prix et al. 2011; Ho et al. 2020a; Yim & Jones 2020), and gravitational-wave searches at glitch epochs

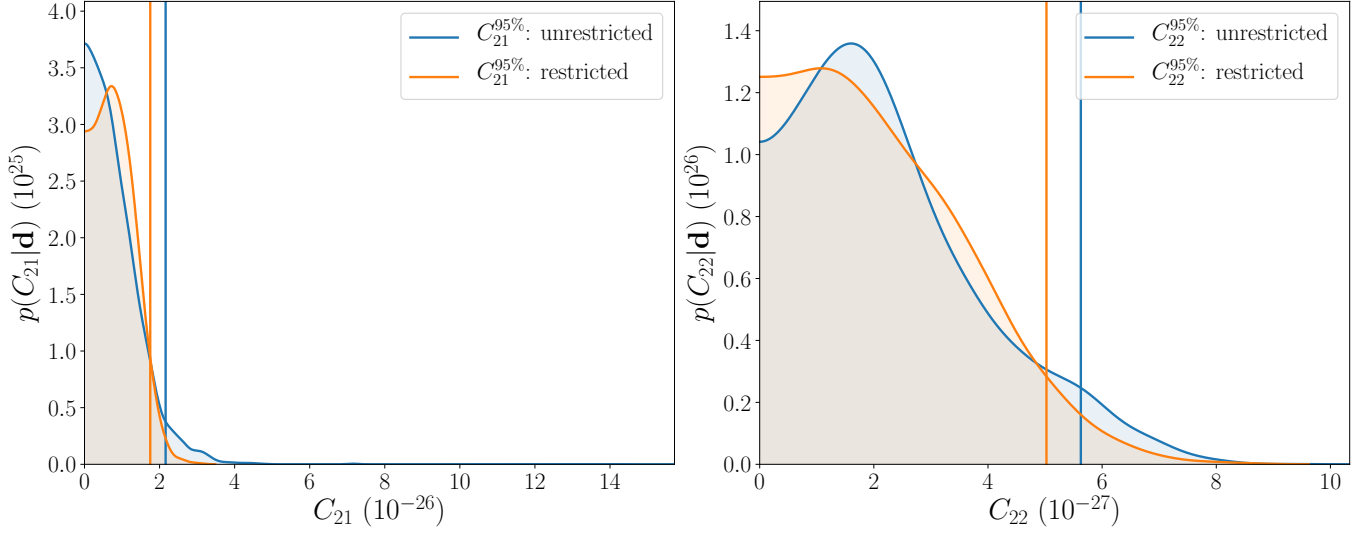


Figure 5. Posterior probability distributions for the amplitudes C_{21} and C_{22} with unrestricted and restricted priors on the pulsar orientation. The 95% credible upper limits are shown as vertical colored lines.

of other pulsars have been conducted (Keitel et al. 2019). It is therefore vital to continue to monitor the spin evolution of PSR J0537–6910, not only to obtain the timing ephemeris and measure braking indices, but also to know when this pulsar undergoes a glitch. Since the spin period of PSR J0537–6910 is only detectable at X-ray energies, *NICER* is the only effective means to perform the necessary observations. Fortunately *NICER* is anticipated to operate until at least late 2022, overlapping with the fourth observing run of LIGO/Virgo and KAGRA (Aso et al. 2013), which is likely to begin in 2022 and continue into 2023.

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Facility: NICER

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